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THE INITIAL-FINAL MASS RELATIONSHIP: SPECTROSCOPY OF WHITE DWARFS IN NGC 2099 (M37)¹

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ABSTRACT

We present new observations of very faint white dwarfs (WDs) in the rich open star cluster NGC 2099 (M37). Following deep, wide-field imaging of the cluster using the Canada-France-Hawaii Telescope, we have now obtained spectroscopic observations of candidate WDs using both the Gemini Multi-Object Spectrograph on Gemini North and the Low-Resolution Imaging Spectrometer on Keck. Of our 24 WD candidates (all fainter than $V = 22.4$), 21 are spectroscopically confirmed to be bona fide WDs, four or five of which are most likely field objects. Fitting 18 of the 21 WD spectra with model atmospheres, we find that most WDs in this cluster are quite massive ($0.7\text{--}0.9 M_{\odot}$), as expected given the cluster's young age (650 Myr) and, hence, high turnoff mass ($\sim 2.4 M_{\odot}$). We determine a new initial-final mass relationship and almost double the number of existing data points from previous studies. The results indicate that stars with initial masses between 2.8 and $3.4 M_{\odot}$ lose 70%–75% of their mass through stellar evolution. For the first time, we find some evidence of a metallicity dependence on the initial-final mass relationship.

Subject headings: open clusters and associations: individual (NGC 2099) — techniques: spectroscopic — white dwarfs

1. INTRODUCTION

The initial-final mass relationship connects the mass of the final products of stellar evolution for intermediate-mass stars, i.e., white dwarfs (WDs), to their progenitor mass. It is a required input for the determination of the ages and distances of globular clusters from modeling their WD cooling sequences (Hansen et al. 2004), for constraining chemical evolution in galaxies, for determining supernova rates (van den Bergh & Tammann 1991), and for understanding feedback processes and star formation in galaxies (e.g., Somerville & Primack 1999). Yet, despite its fundamental importance, this relation remains poorly constrained observationally.

The first attempt to map the relation was made by Weidemann (1977) by comparing theoretical models (e.g., Fusi-Peccini & Renzini 1976) of mass loss to the masses of a few WDs in both the Pleiades and Hyades star clusters. Since then most of the work has focused on using observations of WDs in young open star clusters to provide empirical constraints on the relationship (Koester & Reimers 1981, 1985, 1993, 1996; Reimers & Koester 1982, 1989, 1994; Weidemann & Koester

1983; Weidemann 1987; Jeffries 1997). This effort, spanning almost two decades, is summarized in Weidemann (2000). The result is a monotonically increasing relationship based on about 20 data points, from observations of about a dozen star clusters.

A few recent studies, such as Claver et al. (2001) and Williams et al. (2004), have been successful in finding a half-dozen WDs each in the Praesepe and NGC 2168 clusters. The synthesis of these data with the earlier studies results in an initial-final mass relationship displaying a fair amount of scatter. For example, depending on whose data one uses, initial mass stars of $3\text{--}4 M_{\odot}$ can produce WDs ranging anywhere from 0.65 to $0.8 M_{\odot}$. This result suggests that the relationship may have a stochastic component.

It is desirable to find a young, rich star cluster with a large number of WDs that can be spectroscopically studied. NGC 2099 is such a cluster, with 50 WD cluster candidates measured using imaging observations (Kalirai et al. 2001a). The cluster is very rich, containing over 4000 stars, and has a main-sequence turnoff of $\sim 2.4 M_{\odot}$. Its distance modulus is $(m - M)_V = 11.5$, and its reddening is $E(B - V) = 0.23$. In this Letter, we present a new WD initial-final mass relationship based on 18 WDs in NGC 2099. In the companion Letter, we address the surprising result that all WDs in NGC 2099 (as well as those found in other young open clusters) are all DA spectral type (Kalirai et al. 2005).

2. OBSERVATIONS

Imaging and spectroscopic observations of NGC 2099 were obtained with the Canada-France-Hawaii, Gemini North, and Keck I telescopes. In our wide-field Canada-France-Hawaii Telescope (CFHT) imaging study (Kalirai et al. 2001a), we found ~ 67 WD candidates in the central $15'$ of NGC 2099. Based on comparisons between the cluster field and a blank field surrounding the cluster, we estimate the rate of field star contamination among our candidates to be $\sim 25\%$ (thus yielding 50 cluster WD candidates). We then plotted the locations of all of these faint-blue objects on the sky and obtained further observations of three smaller subfields in the cluster. These fields were chosen to maximize the number of WD candidates

¹ Based on observations with Gemini (run ID GN-2002B-Q-11) and Keck. Gemini is an international partnership managed by the Association of Universities for Research in Astronomy, Inc., under a cooperative agreement with the National Science Foundation. The W. M. Keck Observatory, which is operated as a scientific partnership among the California Institute of Technology, the University of California, and NASA, was made possible by the generous financial support of the W. M. Keck Foundation.

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(all objects in the faint-blue end of the color-magnitude diagram [CMD] were treated as WD candidates). With Gemini, we imaged three $5'.5 \times 5'.5$ fields using the Gemini Multi-Object Spectrograph (GMOS; Murowinski et al. 2003). With Keck, we imaged the same three fields with the Low-Resolution Imaging Spectrometer (LRIS) that has a $5' \times 7'$ field of view (Oke et al. 1995). These imaging data were not significantly deeper than the original CFHT data and were only used to ensure astrometric accuracy for the spectroscopy.

Multiobject spectra were obtained for a single Gemini field and for two of the Keck fields. The Gemini observations used the B600 grating, which simultaneously covers 2760 \AA (centered at $\sim 4700 \text{ \AA}$). The data were binned by a factor of 4 in the spectral direction to improve the signal-to-noise ratio (S/N). The Keck observations used the 600/4000 grism (blue side), which simultaneously covers 2580 \AA (centered at 4590 \AA). For the Gemini field, we obtained 22 individual 1 hr exposures spread over 22 days, taken mostly at low air masses (<1.2) and good seeing ($\sim 0''.8$). For the Keck fields, we obtained four 2000 s exposures in each of the two fields, also at subarcsecond seeing.

The Gemini spectroscopic data were reduced using the Gemini IRAF package, version 1.3. The Keck data were reduced using a set of python routines written by D. Kelson (2004, private communication; Kelson 2003). The individual exposures were bias-subtracted, flat-fielded, cleaned for cosmic rays, wavelength-calibrated, sky-subtracted, extracted, combined, and flux-calibrated (using bright standard stars) within each of these programs. Details of the steps involved in each of these procedures will be provided in a forthcoming paper (J. S. Kalirai et al. 2005, in preparation). The only major problem occurred for some of the Keck field 2 spectra, which were taken at high air masses, and so the bluest flux was lost as a result of atmospheric dispersion. Fortunately, two of the three stars that we were fitting from this field have also been observed in the higher S/N Gemini data.

In total, we obtained spectroscopy of 24 individual WD candidates in the field of NGC 2099 (three of these stars turned out not to be WDs). Therefore, despite sampling only 14% of the total cluster area, we include almost one-third of the total WD population (cluster and field) given the careful positioning of the fields. This is therefore the largest individual star cluster WD sample that has ever been spectroscopically acquired. A CMD showing the locations of the NGC 2099 WDs is shown in Figure 1 (based on CFHT data). The 18 circles represent those objects that we were able to spectroscopically fit (see next section). Four of these, shown as open circles, represent those objects that have inconsistent theoretical magnitudes (from fitting the spectra) as compared to the observed magnitudes, assuming they are cluster members. Three WDs, which we could not fit with models, are shown as crosses. All of the complete spectra are shown in the companion Letter (Kalirai et al. 2005).

3. ANALYSIS

Using the techniques described in Bergeron et al. (1992), we determine T_{eff} and $\log g$ for each WD. The line profiles are first normalized using two points on the continuum on either side of each absorption line. Therefore, the fit should not be affected by the flux calibration unless there is a strange “kink” or slope change at the location of a Balmer line. The fitting of the line shapes uses the nonlinear least-squares method of Levenberg-Marquardt (Press et al. 1986). The χ^2 statistic is calculated and minimized for combinations of T_{eff} and $\log g$, using normalized model line profiles of all absorption lines

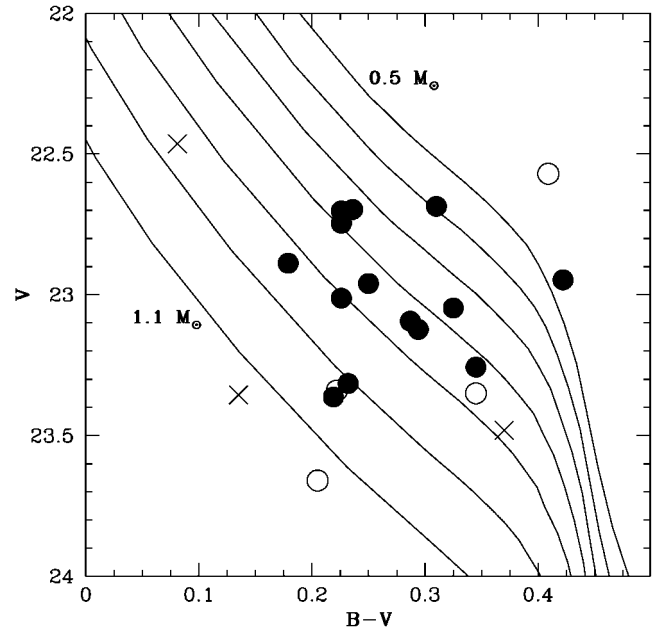


FIG. 1.—Faint-blue end of the CMD of NGC 2099 shown with all spectroscopically confirmed WDs in this work (21 objects, based on CFHT photometry). The 14 filled circles are those with consistent theoretical and observed magnitudes (assuming cluster membership), the four open circles have inconsistent theoretical magnitudes, and the three crosses are objects that we could not fit to spectral models. The model sequences represent WD cooling models (Fontaine et al. 2001) at increments of $0.1 M_{\odot}$, with $0.5 M_{\odot}$ at the top and $1.1 M_{\odot}$ at the bottom.

simultaneously. The resulting 1σ errors in T_{eff} and $\log g$ were tested by simulating synthetic spectra with the same number of absorption lines, and similar S/Ns, and measuring the output parameters from fitting these spectra. These results are found to have errors slightly less than those in the true spectra (as expected given the small errors in flux calibration and other defects), and so we use the true spectra errors. Masses (M_i) and WD cooling ages (t_{cool}) are found by using the updated evolutionary models of Fontaine et al. (2001) for thick hydrogen layers [$q(\text{H}) = M_{\text{H}}/M = 10^{-4}$] and helium layers of $q(\text{He}) = 10^{-2}$. The core is assumed to be a 50/50 C/O mix. In Figure 2 we present the model atmosphere fits for each WD.

The WD cooling age represents the time that each of these stars has spent traversing from the tip of the asymptotic giant branch (AGB) down to its present WD luminosity. We can now calculate the progenitor main-sequence lifetime (t_{ms} , the total lifetime of the star up to the tip of the AGB) assuming an age for the cluster. In Kalirai et al. (2001a), we fitted the NGC 2099 main-sequence to solar metallicity isochrones and determined an age of 520 Myr. Recently, C. Deliyannis et al. (2004, private communication) have spectroscopically measured the cluster metallicity to be subsolar ($Z = 0.011 \pm 0.001$) and the reddening to be $E(B - V) = 0.23 \pm 0.01$, slightly larger than the value that we used. Using these parameters, the age of NGC 2099 is now calculated to be 650 Myr using the same Ventura et al. (1998) models and procedure as described in Kalirai et al. (2001a). We also derive similar ages using the Padova group (620 Myr; Girardi et al. 2000) and Yale-Yonsei isochrones (630 Myr; Yi et al. 2001) for this metallicity (J. S. Kalirai et al. 2005, in preparation).

The t_{ms} for each star is determined by subtracting the WD cooling ages (t_{cool}) from the cluster age (650 Myr). The initial progenitor masses (M_i) for the WDs are then calculated using

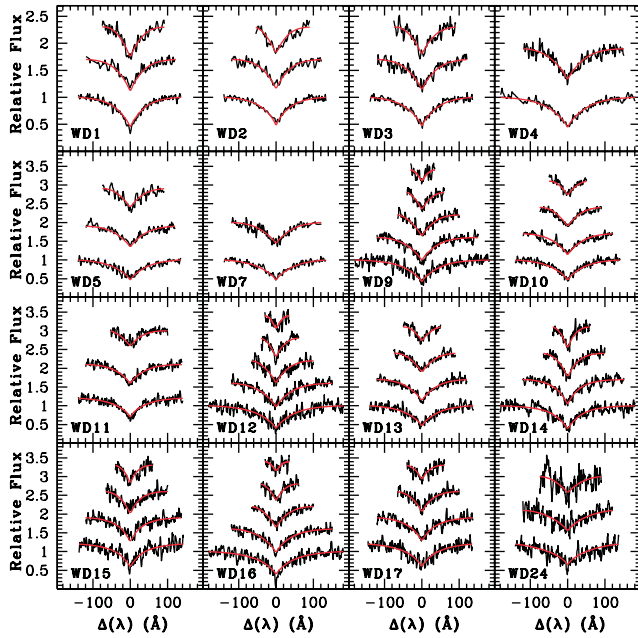


Fig. 2.—Spectroscopic fits (red lines) shown for the hydrogen Balmer lines of 16 WDs in NGC 2099. Within each panel, the lines are H β (bottom), H γ , H δ , He, and H8 (top). The number of higher order lines varies depending on the spectral coverage of each star. The first six objects (WD1–WD7) have the highest S/Ns and were obtained using Gemini. The remaining stars (Keck objects) have lower S/Ns but have bluer spectral coverage (hence the fitting of higher order lines). The identifications are the same as those in Kalirai et al. 2005, in which we present the full spectrum for all WDs. The corresponding T_{eff} , $\log g$, and M_f of these WDs are given in Table 1. The masses of two additional objects are measured as discussed in the text.

the Hurley et al. (2000) models for $Z = 0.01$. For this metallicity, these models give very similar values to those derived from the Ventura et al. (1998) models. For these masses, a 100 Myr cluster age difference would only cause an $\sim 0.3 M_{\odot}$ initial mass difference (for 350 Myr lifetimes). This is a very small effect on

our initial-final mass relationship, and therefore the results are not highly sensitive to the derived cluster age. Table 1 summarizes the derived parameters for each star, with 1σ error bars. Also given are the theoretical magnitudes (from fitting the spectra) and the observed magnitudes [assuming $(m - M)_v = 11.5$ to NGC 2099].

4. THE INITIAL-FINAL MASS RELATIONSHIP

In Figure 3 (top, filled circles) we present the initial-final mass relationship for the 16 WDs in Figure 2 as well as two other objects, WD6 and WD21 (see Kalirai et al. 2005 for the spectra of these objects). For these two stars, we could not determine an accurate $\log g$ as a result of the spectra having too low S/Ns. However, the effective temperatures are well constrained, and therefore we derive the masses of the stars by combining this information with the luminosities (assuming they are cluster members). This gives the radius, which, coupled with a mass-radius relation, gives the mass of the stars. The masses of these two stars, as well as their progenitor masses, are found to be in good agreement with others in the cluster (WD6 has $M_f = 0.92$ and $M_i = 3.25$, and WD21 has $M_f = 0.85$ and $M_i = 3.13$). Two of the WDs shown in Table 1 have negative progenitor lifetimes. This is due to the WD cooling age of these stars being larger than the cluster age. For WD17, we artificially set its initial mass to $7 M_{\odot}$ (the most massive star that produces a WD in the Ventura et al. 1998 models). For WD15, we compute a 95% confidence lower limit of 408 Myr for t_{cool} and determine the initial mass based on this cooling age. These stars are both plotted with open circles and an arrow pointing to higher masses to reflect the lower limits (Fig. 3, top). It is unlikely that these objects are field WDs given their unusually high masses (see Madej et al. 2004 for field WD mass distribution).

Figure 3 (top) shows that, with just one star cluster, we have nearly doubled the number of data points on the initial-final mass plane. This is remarkable considering the time and effort required to establish the previous constraints (shown as triangles). Furthermore, half of our data points sit along a very

TABLE 1
DERIVED PARAMETERS OF WDs

Star	α (J2000)	δ (J2000)	T_{eff} (K)	$\log g$	M_f (M_{\odot})	M_v (Theory)	M_v (Obs.)	$B - V$ (Obs.)	t_{cool} (Myr)	t_{ms} (Myr)	M_i (M_{\odot})
WD1	05 52 26.46	32 30 47.9	16500 ± 800	7.75 ± 0.19	0.48 ± 0.09	10.72	11.07	0.41	111 ± 32	539 ± 32	$2.79^{+0.05}_{-0.06}$
WD2	05 52 17.16	32 29 04.1	19900 ± 900	8.11 ± 0.16	0.69 ± 0.10	10.90	11.20	0.23	108 ± 36	542 ± 36	$2.76^{+0.08}_{-0.05}$
WD3	05 52 36.25	32 32 51.5	18300 ± 900	8.23 ± 0.21	0.76 ± 0.13	11.21	11.39	0.18	179 ± 64	471 ± 64	$2.92^{+0.16}_{-0.14}$
WD4	05 52 34.53	32 29 12.4	16900 ± 1100	8.40 ± 0.26	0.87 ± 0.15	11.65	11.51	0.23	329 ± 133	321 ± 133	$3.36^{+0.76}_{-0.40}$
WD5	05 52 25.11	32 30 55.7	18300 ± 1000	8.33 ± 0.22	0.83 ± 0.14	11.39	11.62	0.29	224 ± 85	426 ± 85	$3.02^{+0.26}_{-0.19}$
WD6 ^a	05 52 18.08	32 29 38.4	17200 ± 1600	...	0.92 ± 0.15	11.76	11.76	0.35	300 ± 100	350 ± 100	$3.25^{+0.44}_{-0.28}$
WD7	05 52 36.96	32 33 29.9	17800 ± 1400	8.42 ± 0.32	0.88 ± 0.19	11.58	11.82	0.23	303 ± 148	347 ± 148	$3.26^{+0.78}_{-0.40}$
WD9	05 52 25.14	32 40 03.6	15300 ± 400	8.00 ± 0.08	0.61 ± 0.05	11.20	11.19	0.31	202 ± 28	448 ± 28	$2.97^{+0.06}_{-0.07}$
WD10	05 52 25.82	32 36 03.8	19300 ± 400	8.20 ± 0.07	0.74 ± 0.04	11.08	11.20	0.24	131 ± 19	519 ± 19	$2.81^{+0.04}_{-0.02}$
WD11	05 52 26.57	32 35 25.4	23000 ± 600	8.54 ± 0.10	0.96 ± 0.06	11.34	11.25	0.23	153 ± 30	498 ± 30	$2.86^{+0.07}_{-0.07}$
WD12	05 52 11.96	32 40 39.5	13300 ± 1000	7.91 ± 0.12	0.55 ± 0.07	11.34	11.45	0.42	313 ± 70	337 ± 70	$3.30^{+0.30}_{-0.22}$
WD13	05 52 16.21	32 38 18.5	18200 ± 400	8.27 ± 0.08	0.79 ± 0.05	11.30	11.46	0.25	183 ± 28	467 ± 28	$2.93^{+0.06}_{-0.07}$
WD14	05 52 06.43	32 38 29.4	11400 ± 200	7.73 ± 0.16	0.45 ± 0.08	11.38	11.55	0.33	318 ± 61	332 ± 61	$3.31^{+0.26}_{-0.20}$
WD15 ^b	05 52 27.42	32 35 50.0	11300 ± 200	8.35 ± 0.15	0.83 ± 0.09	12.30	11.85	0.35	786 ± 189	...	>3.73
WD16	05 52 22.89	32 36 29.4	13100 ± 500	8.34 ± 0.10	0.83 ± 0.06	11.97	11.86	0.22	542 ± 94	108 ± 94	$5.20^{+0.27}_{-0.20}$
WD17 ^b	05 52 29.29	32 37 06.7	12300 ± 400	8.72 ± 0.10	1.05 ± 0.05	12.77	12.16	0.21	1346 ± 256	...	>7
WD21 ^a	05 52 26.42	32 30 04.3	17200 ± 2000	...	0.85 ± 0.15	11.59	11.59	0.29	261 ± 100	389 ± 100	$3.13^{+0.36}_{-0.25}$
WD24	05 52 25.34	32 29 35.2	18700 ± 1100	8.82 ± 0.24	1.11 ± 0.13	12.24	11.84	0.22	492 ± 237	158 ± 237	$4.43^{+0.27}_{-0.32}$

NOTE.—Units of right ascension are hours, minutes, and seconds, and units of declination are degrees, arcminutes, and arcseconds.

^a Masses for these two WDs are determined by assuming that they are cluster members (see § 4).

^b The cooling time for these two WDs (t_{cool}) is larger than the cluster age (see § 4).

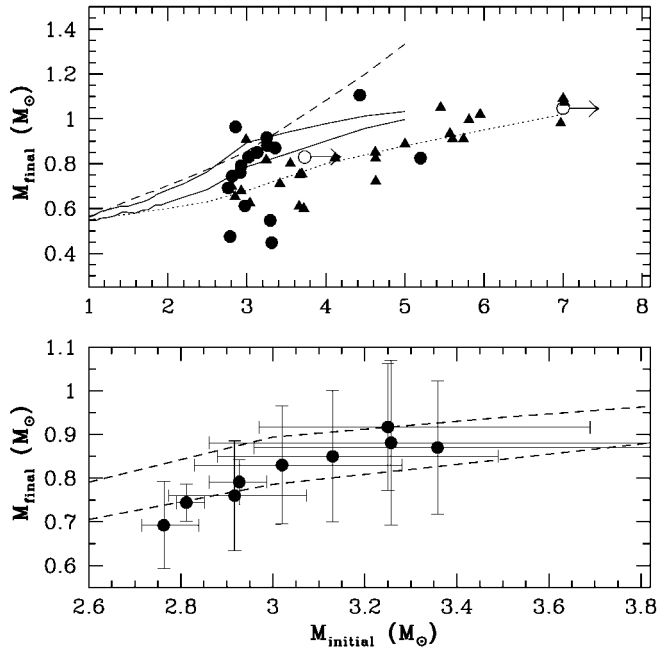


FIG. 3.—*Top*: WD initial-final mass relationship shown for the 18 WDs spectroscopically fitted in this work (*circles*) and all previous constraints (*triangles*). Also shown are several semiempirical and theoretical relations as discussed in the text. *Bottom*: Closer look at those WDs that form the tight sequence with $M_i = 2.8\text{--}3.4 M_\odot$ and are bound by the two Marigo (2001) theoretical relations.

tight sequence with initial masses in the range $2.8\text{--}3.4 M_\odot$ and final masses in the range $0.7\text{--}0.9 M_\odot$. Fitting these data points and their two-dimensional errors to a straight line, we determine a slope of 0.33 and an intercept of -0.19 . The residual rms of the fit is 0.03. This suggests that stars with masses between 2.8 and $3.4 M_\odot$, and metallicities slightly less than solar, will lose 70%–75% of their mass through stellar evolution.

Four of our stars are seen well below the tight sequence of points discussed above. Although only one of these four stars has an inconsistent theoretical magnitude, they could still be field WDs located near the cluster. Not only are their masses consistent with the distribution of field WDs ($M = 0.56 M_\odot$; Sloan Digital Sky Survey data; Madej et al. 2004), but we also expect four or five WDs in our sample to be field stars based on WD number counts in blank fields surrounding NGC 2099 (Kalirai et al. 2001a). The lower masses of these stars could also be attributed to binary star evolution and mass transfer. We will investigate these possibilities further in J. S. Kalirai et al. (2005, in preparation).

We have plotted several theoretical and semiempirical relations in Figure 3. The dotted line is the revised semiempirical initial-final mass relationship from Weidemann (2000). The

dashed line is the theoretical initial-final mass relationship from the Padova group stellar evolution models (Girardi et al. 2000; L. Girardi 2004, private communication) for $Z = 0.008$. The two solid lines are theoretical initial-final mass relations from Marigo (2001), for $Z = 0.008$ (*top*) and $Z = 0.02$ (*bottom*). The latter models have improved mass-loss mechanisms in their post-main-sequence evolutionary phases and therefore should be preferred over the others (L. Girardi 2004, private communication). Considering this, it is interesting to note that half of our NGC 2099 cluster WDs reside in a region of the initial-final mass relationship bound by the two Marigo (2001) relations (see bottom panel of Fig. 3). This is what we would expect, given that the cluster metallicity falls between that used in the two models ($Z = 0.011$). We point out, however, that the errors in our measurements are too large for us to be able to discriminate between the two relations. Our data points are, however, mostly located above the other previous constraints. This can also be explained by the fact that almost all of the other clusters are solar or higher metallicity, resulting in more efficient mass loss during stellar evolution (Marigo 2001). Although this could be a systematic effect (e.g., the masses of the WDs in other clusters have been derived using different models), these results may be suggesting that, for the first time, we are seeing the effects of metallicity on the initial-final mass relationship. In J. S. Kalirai et al. (2005, in preparation), we will eliminate these systematics by reevaluating all of the previous data on the initial-final mass relation (including the new results of Williams et al. 2004) using a consistent fitting method and set of evolutionary models.

The spectroscopic measurement of masses of WDs with $V \approx 23$ has not been previously accomplished. The number of targets that are accessible at these magnitudes is several orders of magnitude larger than previously identified (see work on the ongoing CFHT Open Star Cluster Survey; Kalirai et al. 2001a, 2001b, 2001c, 2003). With continuing observations of rich open star clusters, we can envision placing more than 100 data points into this very fundamental relation. By observing both younger and older clusters, the entire initial progenitor mass range can be constrained, and a detailed initial-final mass relation can be produced.

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